Icy moons in the Solar System

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Course overview

- Motivation why do we study icy moons. History of exploration - telescope observations, spacecraft missions. Surface characteristics - composition, age, and morphology.
- Interior structure layered models: from gravity, shape, composition. Hydrosphere structure - H₂O phase diagram, presence of oceans. Preferred models for selected satellites.
- Dynamics of the different planetary layers. Thermal evolution - heat sources, heat transfer. Melting/crystallization, anti-freezers. Implications for the long-term stability of subsurface oceans.
- **4.** Selected applications. Overview of future missions.

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Habitability requirements

- 1. material: C, H, N, O, P, S (~98% of bio molecules on Earth)
- 2. solvent to speed up reactions liquid water, ...?
- 3. energy source to sustain metabolism
- 4. stable environment



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- $\rightarrow\,$ goal of interior structure modeling is to characterize the ocean and the water/rock interface



Interior structure of the Earth

- layered spherical shells (crust, mantle, outer & inner core)
- Iayering inferred using the travel times of seismic waves:
- body waves: longitudinal P-waves, shear S-waves
- velocity different in each layer, no shear waves through outer core
- reflections & refractions
- + other waves (surface waves, free oscillations, ...)





Preliminary Reference Earth Model (*Dziewonski & Anderson*, 1981)

Interior structure of icy moons (*Hussmann* +, 2015)

- +, 2015)
 - \blacktriangleright no measurements by seismic network \rightarrow main clues:
 - 1 radius (size) and mass
 - 2 gravity field
 - 3 rotational state and shape
 - 4 magnetic field
 - 5 surface temperatures and heat flow
 - 6 composition of surface and atmosphere
 - 7 activity at the surface

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- density + Mol \rightarrow simple structural models of interior
- shape can further confirm/reject the model (consistent or not)

size

- direct observations of satellite's surfaces by spacecraft imaging systems and ground-based telescopes → physical size
- radius of Titan's solid surface buried beneath 100s km of atmosphere first revealed by radio occultation performed by Voyager 1

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- radar altimeter (Titan)

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- $\rightarrow\,$ identification of mutual gravitational interactions among satellites or between satellites and the parent planet

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mass & gravitational field (2)

- 1 visual observations of satellite motions (Earth-based / spacecraft)
- $\rightarrow\,$ identification of mutual gravitational interactions among satellites or between satellites and the parent planet
- $2\,$ radio tracking of the spacecraft path during the moon flyby
- $\rightarrow\,$ gravitational pull on spacecraft $\rightarrow\,$ acceleration/decceleration
- $\rightarrow\,$ Doppler shift of radio communication signal recorded by Earth's DSN
- $\rightarrow\,$ Doppler data inversion $\rightarrow\,$ characteristics of gravitational field

Average density

- mass + size \rightarrow average density $\rho = \frac{3M}{4\pi R^3}$
- $\rightarrow\,$ important indicator of composition



- main composition (~abundances of outer solar system nebula)
- rock (including iron): olivine $(Mg^{2+},Fe^{2+})_2SiO_4$, pyroxenes $XY(Si,AI)_2O_6$, serpentines $(Mg,Fe)_3Si_2O_5(OH)_4$, ...
- nonporous ice/water

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- density in each layer ~ constant (neglect of compression effects)

$$M = \frac{4}{3}\pi\rho R^{3} = M_{r} + M_{i} = \rho_{r}V_{r} + \rho_{i}V_{i} = \frac{4}{3}\pi \left[\rho_{r}R_{r}^{3} + \rho_{i}\left(R^{3} - R_{r}^{3}\right)\right]$$
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- water/ice density ρ_i ~1000 kg m⁻³ (well constrained)
- rock density $\rho_r \sim 2500$ (hydrated rock) 8000 kg m⁻³ (pure iron)
- nonhydrated rock $\rho_r \sim 3500 \text{ kg m}^{-3}$ (close to ρ_{Io})

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$$R_r = R \left(\frac{\rho - \rho_i}{\rho_r - \rho_i} \right)^{1/2}$$

rock and ice mass fractions can be determined from the density

assuming complete differentiation into a rock core and an icy mantle

	Europa	Ganymede	Enceladus	Titan
<i>M</i> [10 ²² kg]	4.8	14.8	0.01	13.5
<i>R</i> [km]	1565	2631	252	2575

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assuming complete differentiation into a rock core and an icy mantle

$$\rho = \frac{5M}{4\pi R^3}$$

$$\rho_r \sim 3500, \ \rho_i \sim 1000: \ R_r = R \left(\frac{\rho - \rho_i}{\rho_r - \rho_i}\right)^{1/3}$$

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$$M_i = \frac{4}{3}\pi\rho_i(R^3 - R_r^3), m_i = \frac{M_i}{M}$$

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<i>d</i> i [km]	115	732	105	752
<i>M</i> _i [10 ²² kg]	0.3278	4.7602	0.0054	4.6126
m_i [%]	7	32	54	34

Average density and ice mass fraction (*Hussmann* +, 2015)



 degree of differentiation (complete or homogeneous ice-rock mixture) has to be inferred from spacecraft flybys

Hydrostatic equilibrium and satellite shape (*Schubert* +, 2009; *Hemingway* +, 2018)

- ${\scriptstyle \bullet}$ interior of large planetary bodies: high pressures + low viscosities
- → relaxation to hydrostatic shape = shape of strengthless fluid body (inward gravity acceleration balanced by fluid pressure gradient):

$$dP = -\rho(r)g(r)dr$$

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- selfgravitation alone: sphere
- rotation: centrifugal flattening
- synchronous rotation with parent body (tidal locking): permanent elongation along the tidal axis

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- tidal + rotational deformation
 → satellites relax to a 3-axial ellipsoid (a > b > c)



Hydrostatic equilibrium and satellite shape

(Schubert +, 2009; Hemingway +, 2018)

- satellite shape can provide additional clues on the interior structure
- accurate measurements of shape can provide evidence for (or against) the hydrostatic state obtained by the satellite

$$\frac{b-c}{a-c}=\frac{1}{4}$$

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- hydrostatic shape \rightarrow hydrostatic gravitational field
- ▶ nonhydrostatic shape + hydrostatic gravity field → compensation: gravity anomaly due to shape reduced by internal density anomaly



Gravitational potential

Newton integral

$$V(r,\theta,\phi) = G \int_{\Omega} \frac{dm(r',\theta',\phi')}{|\vec{r}-\vec{r'}|}$$
$$= G \int_{0}^{R} \int_{0}^{\pi} \int_{0}^{2\pi} \frac{\rho(r',\theta',\phi')}{|\vec{r}-\vec{r'}|} (r')^{2} \sin\theta' dr' d\theta' d\phi'$$

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expansion into series of spherical harmonic functions

$$V(r,\theta,\phi) = \frac{GM}{r} \sum_{j=0}^{\infty} \left(\frac{R}{r}\right)^j \sum_{m=0}^j P_{jm}(\cos\theta) \left[C_{jm}\cos m\phi + S_{jm}\sin m\phi\right]$$



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• $P_{jm}(\xi)$: associated Legendre functions: $P_{jm}(\xi) = (1 - \xi^2)^{m/2} \frac{d^m}{d\xi^m} P_j(\xi)$


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- C_{jm}, S_{jm}: harmonic coefficients (j degree, m order); note S_{j0} = 0 ∀j



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- i = 1 position of center of mass wrt to the origin:

 $C_{11} = \frac{x}{R}, S_{11} = \frac{y}{R}, C_{10} = \frac{z}{R}$ specially choose: origin in the center of mass

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j = 2 components of the moment of inertia tensor I:

(

$$C_{20} = -\frac{1}{MR^2} \left[I_{zz} - \frac{I_{xx} + I_{yy}}{2} \right]$$
$$C_{21} = \frac{I_{xz}}{MR^2}, \quad S_{21} = \frac{I_{yz}}{MR^2}$$
$$C_{22} = -\frac{1}{4MR^2} \left(I_{xx} - I_{yy} \right), \quad S_{22} = \frac{I_{xy}}{MR^2}$$

specially choose: coordinate axes = principal axes of $I \rightarrow C_{21} = S_{21} = S_{22} = 0$

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- dynamical polar flattening J₂:

$$J_2 = -C_{20} = \frac{1}{MR^2} \left[C - \frac{A+B}{2} \right]$$

- mainly caused by satellite's rotation
- best determined with a polar flyby

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- equatorial bulge C₂₂:

$$C_{22} = \frac{1}{4MR^2} \bigg(B - A \bigg)$$

- pointing toward the primary due to tidal interaction
- best determined by equatorial flyby

Tidal & rotational disturbing potential

(Schubert +, 2009; Hemingway +, 2018)

- tidal & rotational forces:
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centrifugal disturbing potential at the surface:

$$V_r(\theta,\phi) \sim \frac{1}{3} \omega^2 R^2 P_{20}(\cos\theta)$$

tidal disturbing potential (neglecting terms ~e)

$$V_t(\theta,\phi) \sim \frac{GM_p R^2}{a_p^3} \left[\frac{1}{2} P_{20}(\cos\theta) - \frac{1}{4} P_{22}(\cos\theta) \cos(2\phi) \right]$$

 $(a_p \text{ distance to tide-raising body, } M_p \text{ its mass})$

Tidal & rotational disturbing potential (*Schubert* +, 2009; *Hemingway* +, 2018)

- rotation coefficient q_r (deformation due to rotation)
- equatorial centrigual potential vs gravitational potential at surface
- symmetry axis: the polar (z) axis

$$q_r = \frac{\omega^2 R^2}{GM/R} = \frac{\omega^2 R^3}{GM}$$

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- tidal coefficient q_t (deformation due to tides):
- surface tidal potential vs gravitational potential at surface
- symmetry axis: the sub-primary line (x axis)

$$q_{t} = -3 \frac{GM_{p}R^{2}/(a_{p})^{3}}{GM/R} = -3 \left(\frac{R}{a_{p}}\right)^{3} \frac{M_{p}}{M}$$

• Kepler's 3rd law:

$$\frac{(a_p)^3}{T^2} = \frac{(a_p)^3 n^2}{4\pi^2} = \frac{G(M_p + M)}{4\pi^2} \sim \frac{GM_p}{4\pi^2}$$

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$$q_t = -3\left(\frac{R}{a_p}\right)^3 \frac{M_p}{M} = -3\frac{R^3\omega^2}{GM_p}\frac{M_p}{M} = -3\frac{R^3\omega^2}{GM} = -3q_r$$

- magnitudes of rotational and tidal deformation differ by a factor of 3
- opposite sign: flattening vs bulge

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- magnitudes of rotational and tidal deformation differ by a factor of 3
- opposite sign: flattening vs bulge

	Europa	Ganymede	Enceladus	Titan
<i>R</i> [km]	1565	2631	252	2575
$\omega \ [10^{-5} \ { m s}^{-1}]$	2.0	1.0	5.3	0.5
<i>GM</i> [10 ³ km ³ s ⁻²]	3.2	9.9	6.7	9.0
$q_r \ [10^{-4}]$	5.0	1.9	67.5	0.4
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relation between gravity coefficients

$$J_2 = \frac{10}{3} C_{22}$$

- determination of coefficients C_{jm} and S_{jm} from Doppler data
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- Doppler $\rightarrow C_{22} \rightarrow k_f = \frac{4C_{22}}{a_r} \rightarrow \text{polar moment of inertia } C$
- Radau-Darwin equation:

$$\frac{C}{MR^2} = \frac{2}{3} \left[1 - \frac{2}{5} \left(\frac{4 - k_f}{1 + k_f} \right)^{1/2} \right]$$

• core ($\rho_c = \rho_m$, $R_c = R_m$), hydrosphere / core, mantle, hydrosphere



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- mass

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- constant density $(\rho_c = \rho_m = \rho_h = \rho)$: MoI = $\frac{2}{5}$
- $MoI < \frac{2}{5}$: increase of density with depth (differentiation)

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Results of layered models: Europa

(Anderson +, 1998) - 4 gravity flybys

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 - 3-layered models:
 - R_c uncertain (Fe vs Fe-FeS), could be as large as ~0.5R
 - silicate mantle
 - hydrosphere thickness between 80 and 170 km



Results of layered models: Ganymede

(Anderson +, 1996) - 2 gravity flybys

- $\blacktriangleright~{\rm MoI}$ = 0.3105 \rightarrow strongly differentiated
- among the smallest value in the Solar System (cf. Earth 0.334)
- average density \rightarrow thick hydrosphere
- detected intrinsic magnetic field (*Kivelson* +, 1996) \rightarrow metallic core

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- 3-layered models (more likely):
- metallic core of radius 400-1300 km (Fe vs Fe-FeS composition)
- silicate mantle
- hydrosphere ~800 km



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 - 2-layered models:
 - low core density of ~2400 kg m^{-3}
 - hydrosphere ~60 km
 - compensation depth ~30-40 km: thickness of ice crust?
 - negative mass anomaly in the South polar region regional sea?
- $\rightarrow\,$ regional vs global ocean? cannot be answered with gravity data

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 - if $\rho_h{\sim}1000~{\rm kg}~{\rm m}^{-3}{\rm :}~R_c{\sim}2200~{\rm km},~\rho_c{\sim}2500~{\rm kg}~{\rm m}^3,~d_h{\sim}400~{\rm km}$
 - if ρ_h larger and/or MoI smaller: smaller and denser cores



Interior models based on mass



Interior models based on mass and Mol



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models taking into account the gravity field:

- · Europa and Ganymede: iron cores, slightly thicker hydrospheres
- \blacktriangleright Enceladus and Titan: thinner hydrospheres \rightarrow low core densities
- hydrated silicates
- porous material: Enceladus (Choblet +, 2017)
- organic material: Titan (*Néri* +, in rev.)

Course overview

- Motivation why do we study icy moons. History of exploration - telescope observations, spacecraft missions. Surface characteristics - composition, age, and morphology.
- Interior structure layered models: from gravity, shape, composition. Hydrosphere structure - H₂O phase diagram, presence of oceans. Preferred models for selected satellites.
- Dynamics of the different planetary layers. Thermal evolution - heat sources, heat transfer. Melting/crystallization, anti-freezers. Implications for the long-term stability of subsurface oceans.
- **4.** Selected applications. Overview of future missions.

Presence of deep ocean

- due to small density difference between ice and liquid water, presence of the ocean cannot be inferred from the mass and gravity data
- other evidence for liquid water ocean:
- induced magnetic field
- auroral ovals oscillation
- Schumann resonance
- libration & obliquity
- tidal deformation
- heat flux
- surface activity
- ...

- Faraday: time-varying mg. field accompanied by (time-var) el. field
 - conductor in time-varying mg field: surface eddy currents induce secondary field that reduces primary field in the conductor

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(Khurana +, 2009)

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- uniform primary field, dipolar induced field (same frequency)
- primary + induced field avoids the moon
- elmg induction: detection & characterisation of secondary field
- $\rightarrow\,$ information on location, size, shape, and conductivity



- Jupiter's mg field: dipole tilted by ~9.6° \rightarrow primary oscillating field
- additionally: day/night asymmetry in Jupiter's magnetospheric field
- Saturn's mg field: not inclined wrt rotation axis → moons do not sense a systematic time-periodic field in their rest frame



skin depth: distance over which the primary signal decays to 1/e

$$S = (\omega \mu_0 \sigma/2)^{-1/2}$$

- small: large material conductivity and/or high sounding frequency
- obstacle thickness >S: primary wave reflection \rightarrow induced field

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- Jupiter spin period ~10 h:
- pure ice/water skin depth: ~ $10^6 \text{ km} \gg R_E \rightarrow \text{ no induction}$
- ocean water skin depth: ~60 km
 → significant induction
- three-layered model can be used to model the moon's induction response



Induced magnetic field: Europa and Ganymede

(Khurana +, 1998; Kivelson +, 2000; Kivelson +, 2002; Schilling +, 2007)

Europa

- flybys with sufficiently low altitude required for an adequate signal-to-noise ratio to decipher the induced field
- strong evidence that Europa has a subsurface liquid water ocean
- best fit of Galileo data: ocean thickness < 100 km (cannot rule out thicker ocean)

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Ganymede

- discovery of intrinsic magnetic (Kivelson+, 1996)
- satisfactory fits of Galileo data:
- (a) internal field with dipole and quadrupole terms
- (b) internal permanent dipole
 - + induced magnetic dipole from ocean ${\sim}150$ km deep
 - data did not allow to confirm the presence of an ocean

Ganymede: auroral ovals oscillation (*Saur* +, 2015)

- ▶ auroral emission first observed by HST (Hall +, 1998)
- shape: two circumpolar auroral ovals in N & S polar regions
- location: open-closed field line boundary separates mg field lines starting & ending on Ganymede from field lines connecting to Jupiter
- \blacktriangleright locations controlled by time-variable mg environment \rightarrow oscillations
- ocean: primary field reduced by induced field \rightarrow oscillation reduction



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- HST observations: average oscillation of 2.0° ± 1.3°
- model:
- withouth ocean $5.8^\circ\pm1.3^\circ$
- with ocean $2.2^\circ \pm 1.3^\circ$
- ocean depth ~150–250 km



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- synchronously rotating satellite on eccentric orbit:
- long axis points into empty focus: optical libration (Ψ)
- \rightarrow periodic misalignement (long axis vs satellite-planet line)
- \rightarrow gravitational torque \rightarrow oscillations: physical longitudinal libration ($\gamma)$
 - ${\scriptstyle \bullet}\,$ amplitude of γ depends on satellite's interior structure



(credits: Hemingway +, 2018)

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Interior model	Amplitude of forced libration
Homogeneous ellipsoid	0.032°
2-layer hydrostatic	$0.032^\circ-0.034^\circ$
2-layer hydrostatic, including	$0.032^\circ-0.034^\circ$
"polar sea" and depression	
Ellipsoidal core, global ocean,	0.120°
ellipsoidal shell (23 km)	
(2300, 1000, 850 kg/m ³)	
Measured value	$0.120^\circ\pm0.014^\circ$

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- 2 presence of a global ocean rather than a localized polar sea


(Béghin +, 2012)

- 1 Schumann resonance
- Earth:
- global elmg resonance with extremely low frequency (ELF, 3–60 Hz)
- in the cavity formed between the Earth's surface and the ionosphere
- excited by lightning discharges
- modes: 7.83 Hz (fundamental), 14.3, 20.8, 27.3 and 33.8 Hz



(Béghin +, 2012)

- 1 Schumann resonance
- Titan descent of Huygens probe:
- observed ELF (36 Hz) wave: 2nd harmonic of Schumann resonance
- resonating cavity: between layered ionosphere (up to 150 km height) and lower conductive surface beneath non-conductive ground
- excitation: ionospheric current sources
- lower reflector: water-ammonia ocean ~55-80 km below icy crust



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 - 2 tides
 - eccentricity of Titan's orbit (~2.9%)
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- 2 independent determinations from Cassini gravity data: $k_2 = 0.589 \pm 0.150, \ k_2 = 0.637 \pm 0.224$
- $\rightarrow\,$ some global layer within Titan behaves like a fluid on orbital time scales
- (i) very low viscosity layer (an ocean) beneath an outer ice shell
- (ii) low viscosity deep interior

(Baland +, 2011)

- 1 Schumann resonance (Béghin +, 2012)
- 2 tides (less +, 2012)
- 3 obliquity (axial tilt)
- angle between satellite's rotational and orbital axis (~ angle between its equatorial and orbital plane)



(Baland +, 2011)

- 1 Schumann resonance (Béghin +, 2012)
- 2 tides (less +, 2012)
- 3 obliquity (axial tilt)
- angle between satellite's rotational and orbital axis (~ angle between its equatorial and orbital plane)
- Titan's obliquity: ε~0.3° (Seidelmann +, 2007)
- models: (Baland +, 2011)
- completely solid Titan: $\varepsilon = 0.12^{\circ} \pm 0.02^{\circ}$
- Titan with a liquid ocean: $\varepsilon = 0.32^{\circ} \pm 0.02^{\circ}$
- → another indirect evidence for Titan's subsurface ocean



Structure of the hydrosphere - H_2O phase diagram



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- hydrosphere-rock boundary (HRB):
- Europa, Enceladus: ice I
- Ganymede, Titan: high-pressure (HP) ices

Structure of the hydrosphere - H_2O phase diagram



- hydrosphere-rock boundary (HRB):
- Europa, Enceladus: ice I
- Ganymede, Titan: high-pressure (HP) ices
- $\rightarrow\,$ liquid water ocean sandwiched between two ice layers







- · Europa, Enceladus: ocean in direct contact with rock
- · Ganymede, Titan: high-pressure ice decouples ocean from the rock



- · Europa, Enceladus: ocean in direct contact with rock
- $\rightarrow\,$ great for origin of life $\odot\,$
 - · Ganymede, Titan: high-pressure ice decouples ocean from the rock
- $\rightarrow\,$ not so great for origin of life $\odot\,$

Adding more complexity

So far, we have assumed that:

- densities in particular layers are constant
- pure water/ice

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So far, we have assumed that:

- densities in particular layers are constant
- pure water/ice

In reality:

- density changes with mineralogy that depends on (P,T)
- ▶ salts (NaCl, MgSO₄, ...) and/or ammonia in the subsurface oceans
- $\rightarrow\,$ reduction of melting temperature
- $\rightarrow\,$ change of liquid water buoyancy



Course overview

- Motivation why do we study icy moons. History of exploration - telescope observations, spacecraft missions. Surface characteristics - composition, age, and morphology.
- Interior structure layered models: from gravity, shape, composition. Hydrosphere structure - H₂O phase diagram, presence of oceans. Preferred models for selected satellites.
- Dynamics of the different planetary layers. Thermal evolution - heat sources, heat transfer. Melting/crystallization, anti-freezers. Implications for the long-term stability of subsurface oceans.
- **4.** Selected applications. Overview of future missions.

Preferred interior models for selected satellites



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